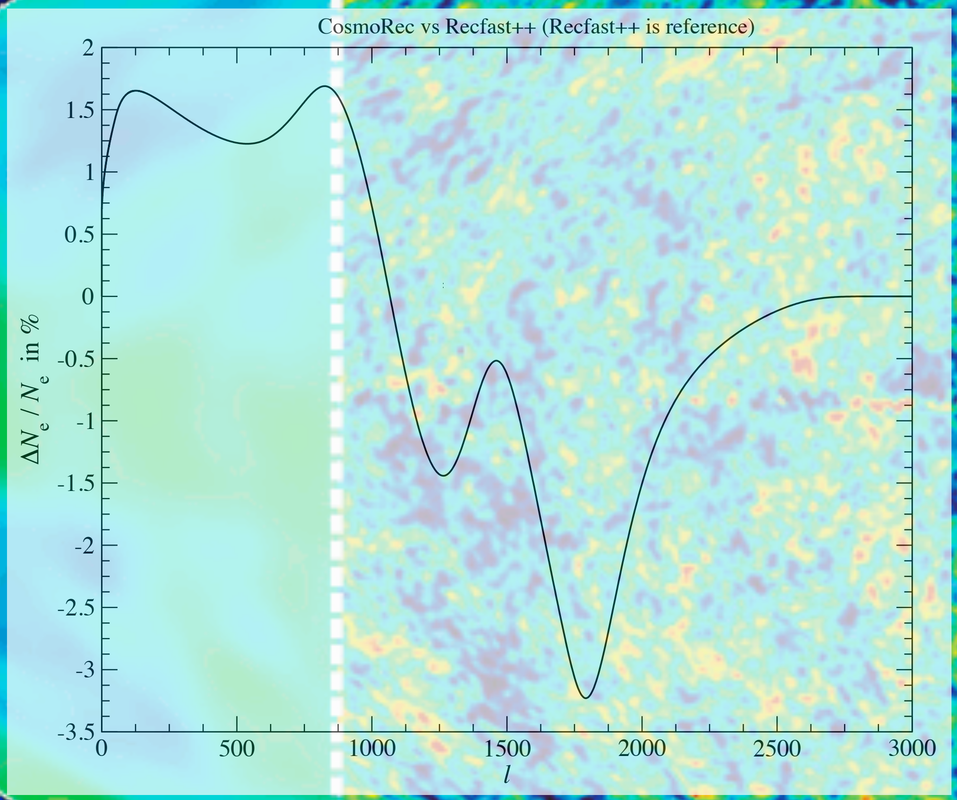
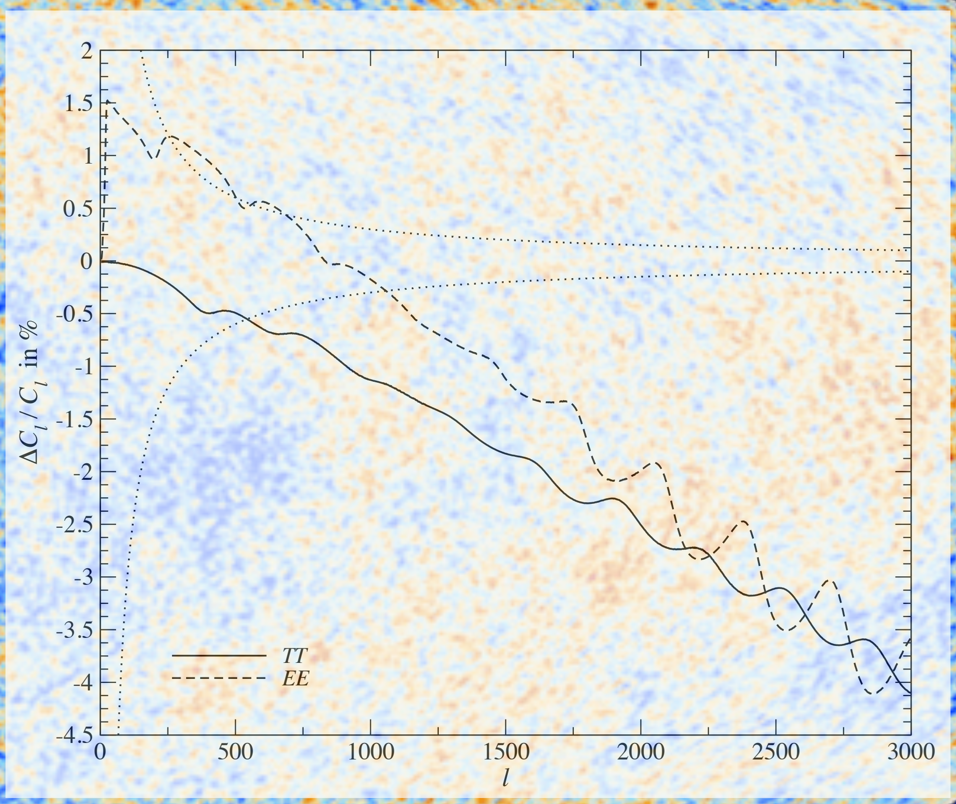
Overview of different recombination codes

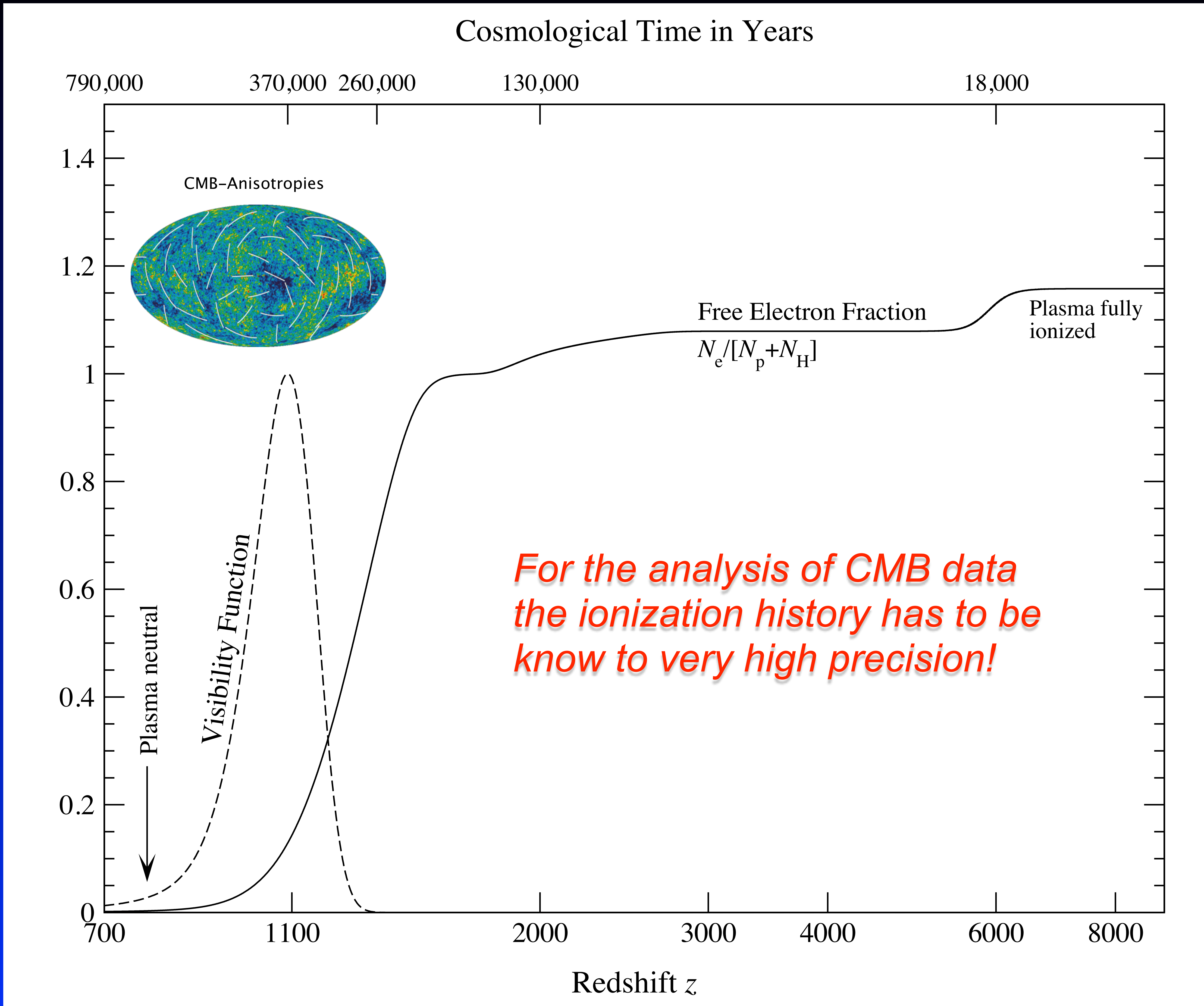




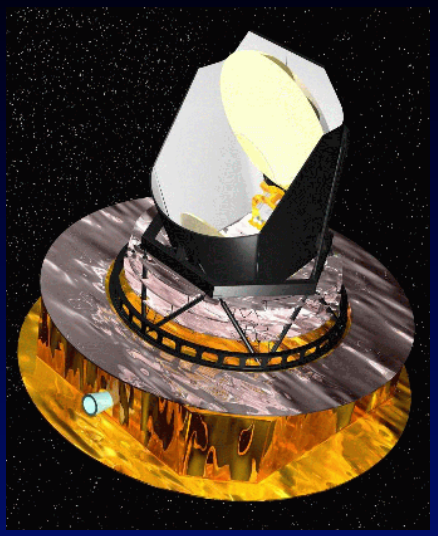
Jens Chluba

School on Cosmological Tools

Madrid, Spain, Nov 12th - 15th, 2013

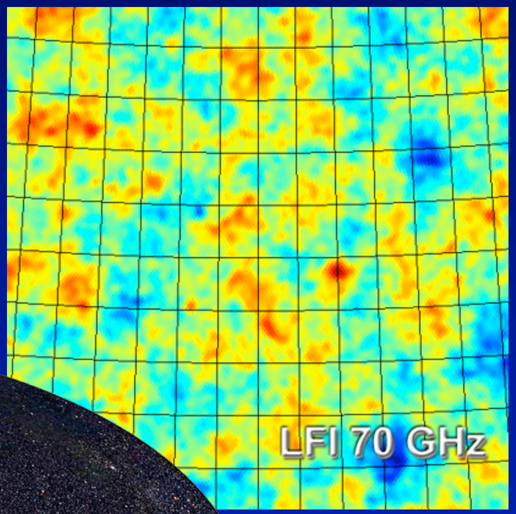

Cosmological Time in Years
790,000
370,000
260,000
130,000
18,000
1.4
1.2
1
0.8
0.6
0.4
0.2
0
CMB–Anisotropies
Free Electron Fraction
$N_e/[N_p+N_H]$
Plasma fully ionized
Visibility Function
Plasma neutral
For the analysis of CMB data the ionization history has to be know to very high precision!
700
1100
2000
3000
4000
6000
8000
Redshift z

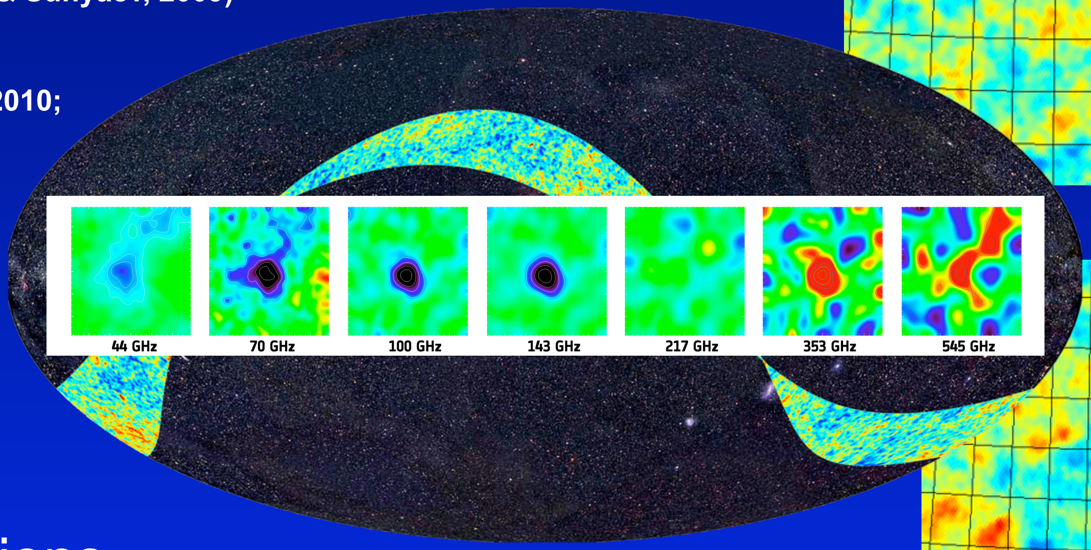
Getting the job done for Planck

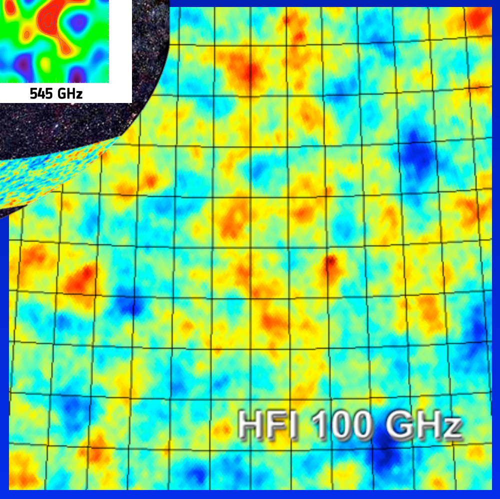


Hydrogen recombination

- Two-photon decays from higher levels
  (Dubrovich & Grachev, 2005, Astr. Lett., 31, 359; Wong & Scott, 2007; JC & Sunyaev, 2007; Hirata, 2008; JC & Sunyaev 2009)
- Induced 2s two-photon decay for hydrogen
  (JC & Sunyaev, 2006, A&A, 446, 39; Hirata 2008)
- Feedback of the Lyman-α distortion on the 1s-2s two-photon absorption rate
  (Kholupenko & Ivanchik, 2006, Astr. Lett.; Fendt et al. 2008; Hirata 2008)
- Non-equilibrium effects in the angular momentum sub-states
  (Rubiño-Martín, JC & Sunyaev, 2006, MNRAS; JC, Rubiño-Martín & Sunyaev, 2007, MNRAS; Grin & Hirata, 2009; JC, Vasil & Dursi, 2010)
- Feedback of Lyman-series photons (Ly[n] → Ly[n-1])
  (JC & Sunyaev, 2007, A&A; Kholupenko et al. 2010; Haimoud, Grin & Hirata, 2010)
- Lyman-α escape problem (*atomic recoil, time-dependence, partial redistribution*)
  (Dubrovich & Grachev, 2008; JC & Sunyaev, 2008; Forbes & Hirata, 2009; JC & Sunyaev, 2009)
- Collisions and Quadrupole lines
  (JC, Rubiño-Martín & Sunyaev, 2007; Grin & Hirata, 2009; JC, Vasil & Dursi, 2010; JC, Fung & Switzer, 2011)
- Raman scattering
  (Hirata 2008; JC & Thomas , 2010; Haimoud & Hirata, 2010)







Helium recombination

- Similar list of processes as for hydrogen
  (Switzer & Hirata, 2007a&b; Hirata & Switzer, 2007)
- Spin forbidden 2p-1s triplet-singlet transitions
  (Dubrovich & Grachev, 2005, Astr. Lett.; Wong & Scott, 2007; Switzer & Hirata, 2007; Kholupenko, Ivanchik&Varshalovich, 2007)
- Hydrogen continuum opacity during He I recombination
  (Switzer & Hirata, 2007; Kholupenko, Ivanchik & Varshalovich, 2007; Rubiño-Martín, JC & Sunyaev, 2007; JC, Fung & Switzer, 2011)
- Detailed feedback of helium photons
  (Switzer & Hirata, 2007a; JC & Sunyaev, 2009, MNRAS; JC, Fung & Switzer, 2011)

$\Delta N_e \,/\, N_e \sim 0.1\ \%$

Recombination code overview

Code	Recfast	Recfast++	CosmoRec
Language	Fortran 77/90 & C	C++	C++
Requirements	-	-	GNU Scientific Lib (GSL)
Solves for	X_p, X_{HeI}, T_e	X_p, X_{HeI}, T_e	X_{1s}, X_{ns}, X_{np}, X_{nd}, T_e
Solves for	X_p, X_{HeI}, T_e	X_p, X_{HeI}, T_e	X_{1s}, X_{ns}, X_{np}, X_{nd}, T_e
ODE-Solver	explicit	implicit (Gears method)	implicit (Gears method)
PDE-Solver	-	-	semi-implicit (Crank-Nicolson)
Approach	derivative fudge	correction function	physics
Simplicity	simple	simpler	pretty big code
Flexibility	limited	better but limited	very flexible
Validity	close to standard cosmology	close to standard cosmology	wide range of cosmologies
Tools	-	ODE Solver	*HI & He Atom, Solvers, Quadrature routines*
Extras	-	DM annihilation	*DM annihilation, high-ν distortion*
Runtime	0.01 sec	0.08 sec	1.5 - 2 sec

Recfast Equations

$$\frac{dx_p}{dz} = \frac{[x_e x_p n_{\rm H}\alpha_{\rm H} - \beta_{\rm H}(1-x_p)e^{-h\nu_{{\rm H}\,2s}/kT_M}][1+K_{\rm H}\Lambda_{\rm H}n_{\rm H}(1-x_p)]}{H(z)(1+z)[1+K_{\rm H}(\Lambda_{\rm H}+\beta_{\rm H})n_{\rm H}(1-x_p)]}, \quad (1)$$

$$\frac{dx_{\rm He\,II}}{dz} = \{[x_{\rm He\,II}x_e n_{\rm H}\alpha_{\rm He\,I} - \beta_{\rm He\,I}(f_{\rm He}-x_{\rm He\,II})e^{-h\nu_{{\rm He\,I}\,2^1s}/kT_M}] \times [1+K_{\rm He\,I}\Lambda_{\rm He}n_{\rm H}(f_{\rm He}-x_{\rm He\,II})e^{-h\nu_{{\rm He\,I}\,2^1p2^1s}/kT_M}]\} \Big/ \{H(z)(1+z)[1+K_{\rm He\,I}(\Lambda_{\rm He}+\beta_{\rm He\,I})n_{\rm H} \times (f_{\rm He}-x_{\rm He\,II})e^{-h\nu_{{\rm He\,I}\,2^1p2^1s}/kT_M}]\}, \quad (2)$$

$$\frac{dT_M}{dz} = \frac{8\sigma_{\rm T}a_R T_R^4}{3H(z)(1+z)m_e c}\frac{x_e}{1+f_{\rm He}+x_e}(T_M-T_R) + \frac{2T_M}{(1+z)}$$

Seager et al, 1999

- Old expressions from Peebles 1969
- second shell quasi-stationary
- recombination rates and escape probabilities fudged
- spin-forbidden transition added to helium equation (Wong, Moss & Scott, 2009)

Recfast

recfast.readme

```
The input interface was designed to look familiar to users of Seljak &
Zaldarriaga's code CMBFAST.  A convenient way to run the program is by using
a file recfast.run of the form:

output.file
Omega_B, Omega_DM, Omega_vac
H_0, T_0, Y_p
Hswitch
Heswitch

For example:
junk.out
0.04 0.20 0.76
70 2.725 0.25
1
6
```

meaning of parameters

write into **`recfast.ini`**

Execute code like **`./recfast < recfast.ini`**

recfast.for

```
c       Modification for H correction (Hswitch):
        write(*,*) 'Modification for H recombination:'
        write(*,*)'0) no change from old Recfast'
    write(*,*)'1) include correction'
        write(*,*)'Enter the choice of modification for H (0-1):'
    read(*,*)Hswitch

C   Fudge factor to approximate the low z out of equilibrium effect
    if (Hswitch .eq. 0) then
      fu=1.14d0
    else
      fu=1.125d0
    end if

C   Modification for HeI recombination (Heswitch):
    write(*,*)'Modification for HeI recombination:'
    write(*,*)'0) no change from old Recfast'
    write(*,*)'1) full expression for escape probability for singlet'
    write(*,*)'   1P-1S transition'
    write(*,*)'2) also including effect of contiuum opacity of H on HeI'
    write(*,*)'   singlet (based in fitting formula suggested by'
    write(*,*)'   Kholupenko, Ivanchik & Varshalovich, 2007)'
    write(*,*)'3) only including recombination through the triplets'
    write(*,*)'4) including 3 and the effect of the contiuum '
    write(*,*)'   (although this is probably negligible)'
    write(*,*)'5) including only 1, 2 and 3'
    write(*,*)'6) including all of 1 to 4'
    write(*,*)'Enter the choice of modification for HeI (0-6):'
    read(*,*)Heswitch
```

Recfast++

Initialization for Recfast++ uses same file as CosmoRec

```
//==================================================================================================
// the above parameters are (default values are given as examples)
//==================================================================================================

2000      == number of redshift points (for the range z= 50-3000 nz=500 is in principle sufficient)
3000      == starting redshift; above z=3400 the Recfast++ Solution should be used.
             This is automatically done in batch mode.
0         == ending redshift; below z=50 the Recfast++ system is solved with rescale dXe/dt

0.24      == Yp
2.725     == T0
0.2678    == Omega_m
0.0444    == Omega_b
0.7322    == Omega_L  (if <=0 it will be computed from the other variables)
0.0       == Omega_k
0.71      == h100
3.04      == N_nu
1.14      == Recfast++ fudge factor (usually leave unchanged)

3         == number of hydrogen shells for ODE problem (currently: 3, 4, 5 or 10; lite only 3)
500       == nS for effective HI rates (nS=10, 20, 50, 100, 128, 200, 300, 400 and 500; lite only 500)
1.0e-24   == dark matter annihilation efficiency in eV/sec (see Chluba 2009).
             Values <= 10^-23 eV/sec are recommended. For larger values the CosmoRec
             calculation breaks down. In Recfast-mode also larger values are possible.

3         == number of helium shells (currently: 2, 3, 5, or 10; lite only 3)
0         == HI absorption during HeI-recombination (0: off; 1: on; 2: on with Diffusion fudge)
0         == spin forbidden transitions for HeI-recombination (0: off; 1: on)
0         == Feedback in Helium levels (positive: no HI abs between the lines
                    negative: with HI abs between the lines)

1     == run PDE part (1) or not (0). In the latter case only ODE system will be solved.
         If this flag is set to 0 only the initial calculation without transfer corrections
         will be performed
2     == correction to 2s-1s channel; 0: no corr; 1: stim. 2s-1s; 2: full correction;
3     == nS for corrections because of two-photon decays.
         If set to <3 then only the diffusion correction is included.
2     == nS for corrections because of Raman-scattering
         If set to <2 then the 1+1 Raman rates are not corrected.

./outputs/ == path for output
.dat       == addition to name of files at the very end

//==================================================================================================
//==================================================================================================
```

./runfiles/parameters.dat

parameters for both Recfast++ & CosmoRec

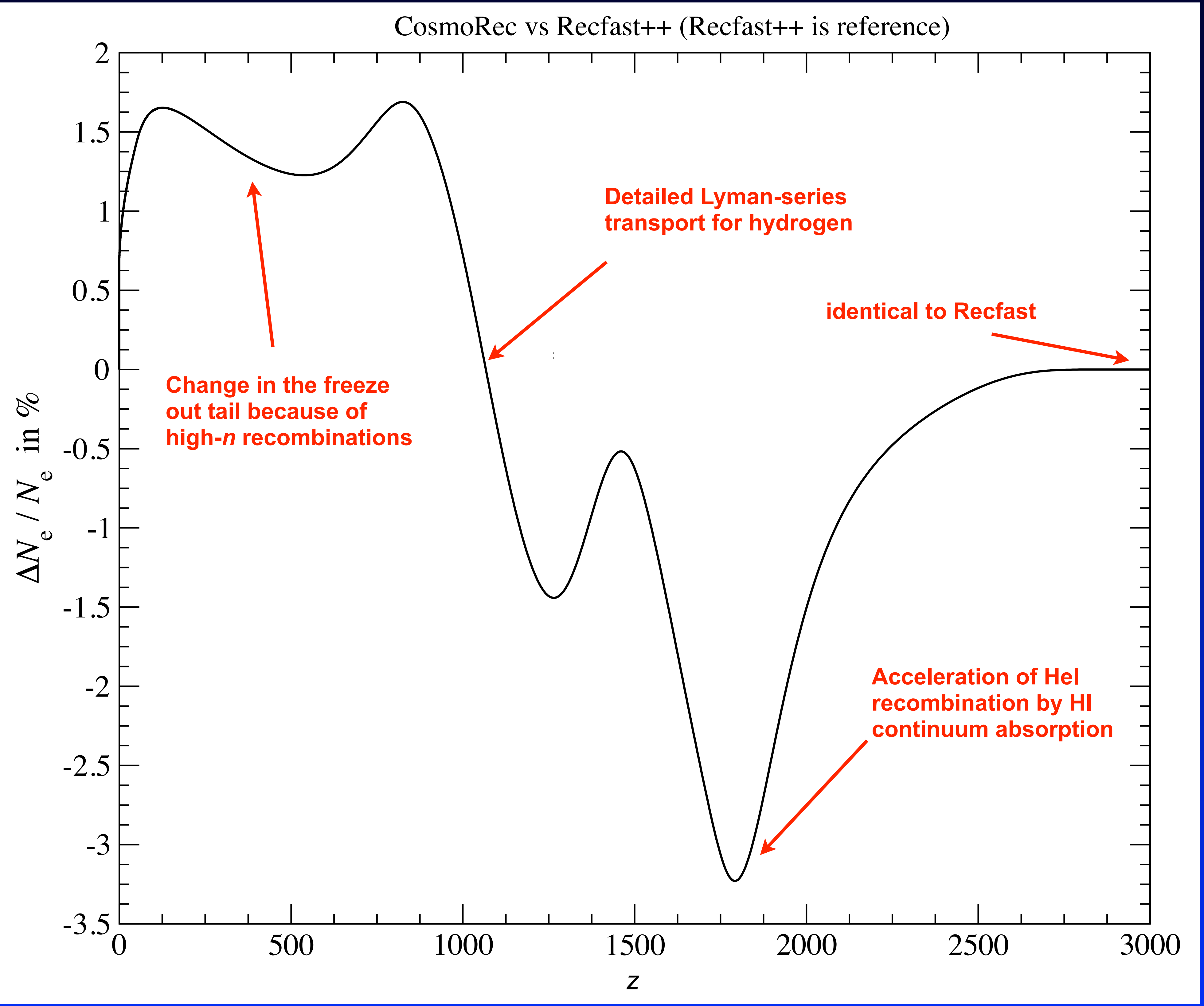
main CosmoRec parameters

Execute Recfast++ like

`./CosmoRec REC runfiles/parameters.dat` **(equivalent to old recfast)**

`./CosmoRec RECcf runfiles/parameters.dat` **(recfast + correction function)**

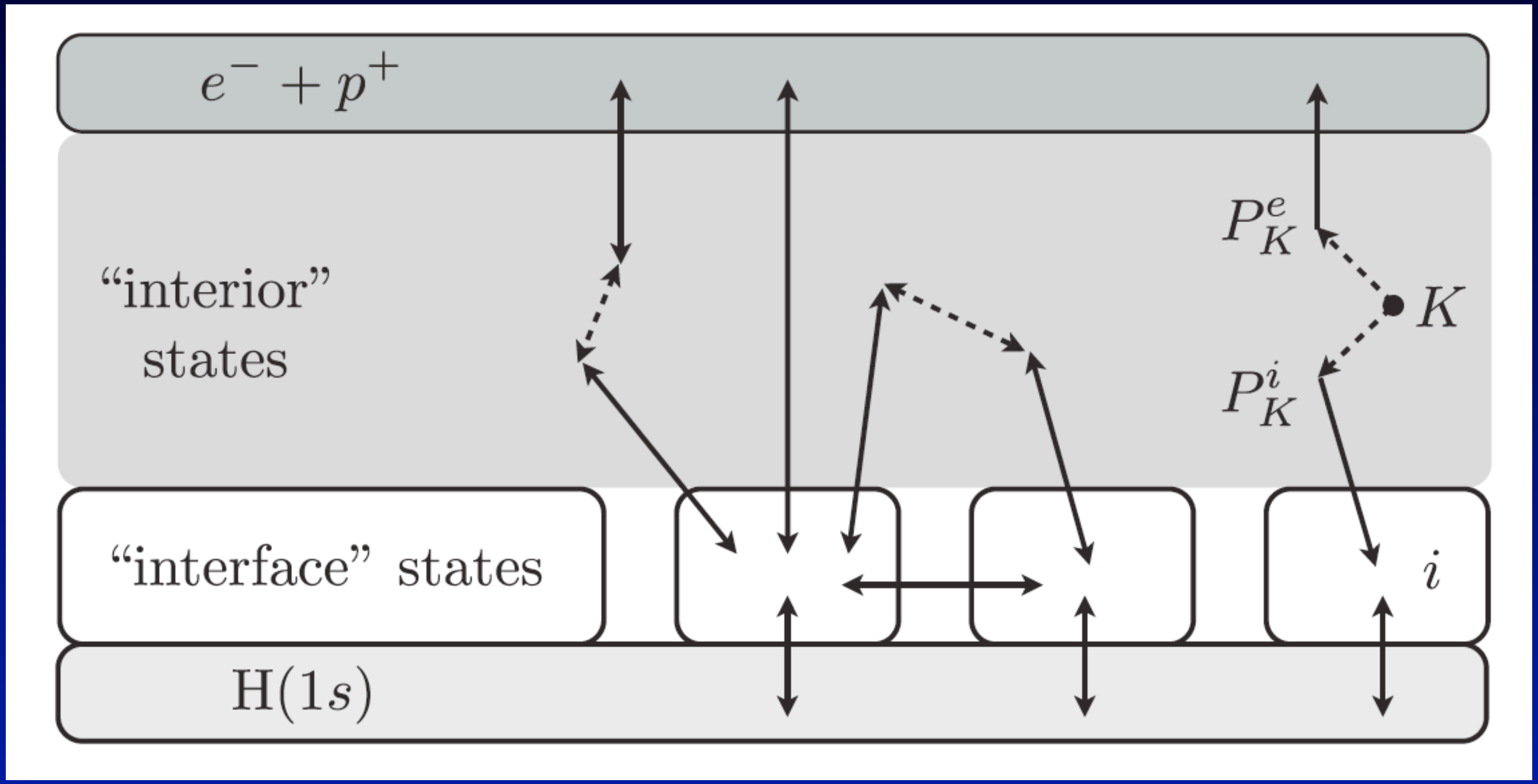
Correction function approach just uses full correction!



Introduced in Rubino-Martin et al, 2009

CosmoRec

Extended Effective Multi-level Atom



CosmoRec & HyRec

- need to treat angular momentum sub-levels separately
- Complexity of problem scales like ~ n^2_{max}
- Full problem pretty demanding (500 shells ≃ 130000 equations!)

$\Longrightarrow$ *effective multi-level approach* (Ali-Haimoud & Hirata, 2010)

- This allowed fast computation of the recombination problem!

CosmoRec parameters

./runfiles/parameters.dat

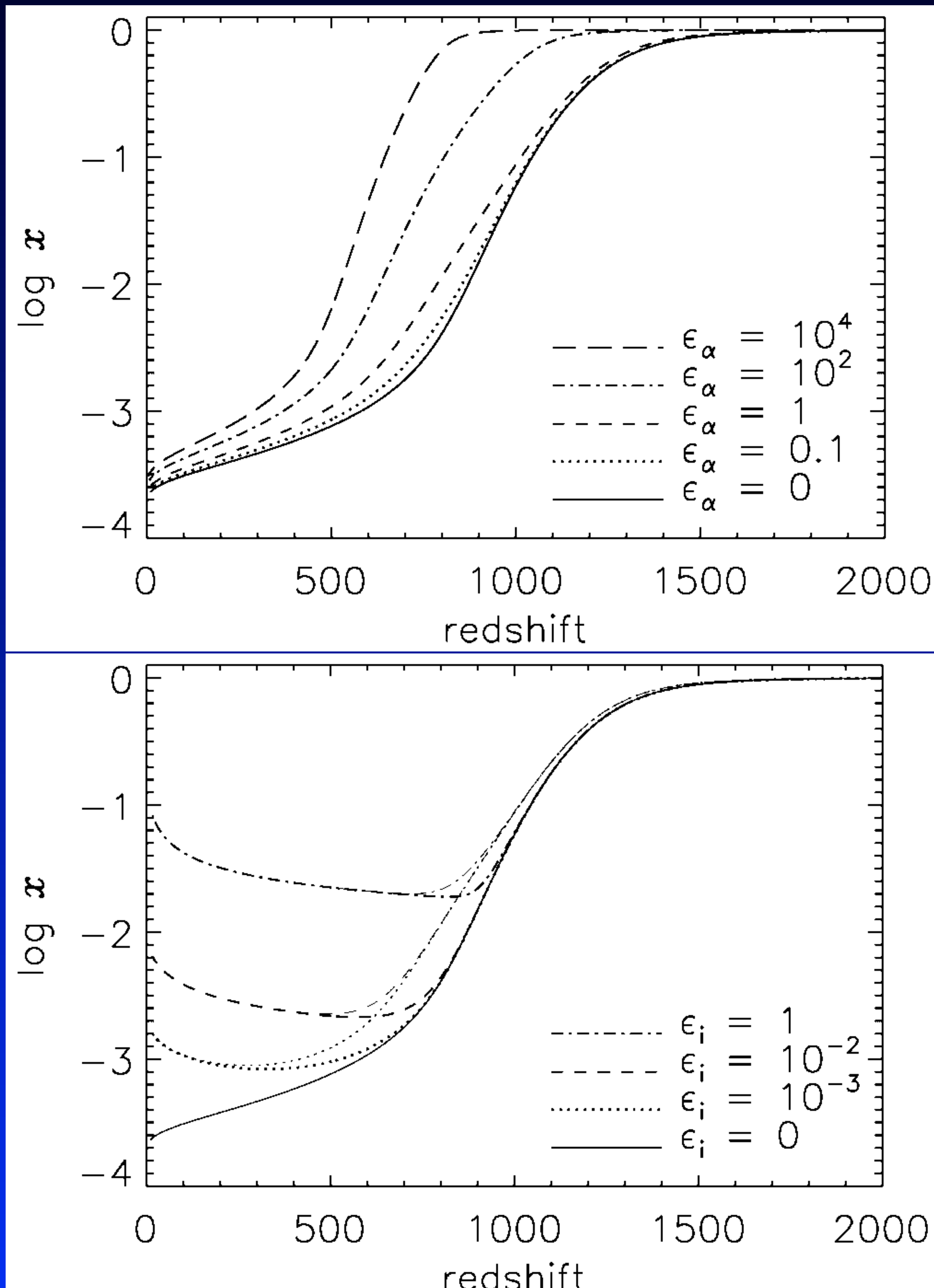
```
3         == number of hydrogen shells for ODE problem (currently: 3, 4, 5 or 10; lite only 3)
500       == nS for effective HI rates (nS=10, 20, 50, 100, 128, 200, 300, 400 and 500; lite only 500)
1.0e-24   == dark matter annihilation efficiency in eV/sec (see Chluba 2009).
             Values <= 10^-23 eV/sec are recommended. For larger values the CosmoRec
             calculation breaks down. In Recfast-mode also larger values are possible.

3         == number of helium shells (currently: 2, 3, 5, or 10; lite only 3)
0         == HI absorption during HeI-recombination (0: off; 1: on; 2: on with Diffusion fudge)
0         == spin forbidden transitions for HeI-recombination (0: off; 1: on)
0         == Feedback in Helium levels (positive: no HI abs between the lines
                    negative: with HI abs between the lines)

1     == run PDE part (1) or not (0). In the latter case only ODE system will be solved.
         If this flag is set to 0 only the initial calculation without transfer corrections
         will be performed
2     == correction to 2s-1s channel; 0: no corr; 1: stim. 2s-1s; 2: full correction;
3     == nS for corrections because of two-photon decays.
         If set to <3 then only the diffusion correction is included.
2     == nS for corrections because of Raman-scattering
         If set to <2 then the 1+1 Raman rates are not corrected.

./outputs/ == path for output
.dat        == addition to name of files at the very end
```

Execute CosmoRec like

```
./CosmoRec runfiles/parameters.dat
```

Annihilation and extra energy release

Extra Sources of Ionizations or Excitations

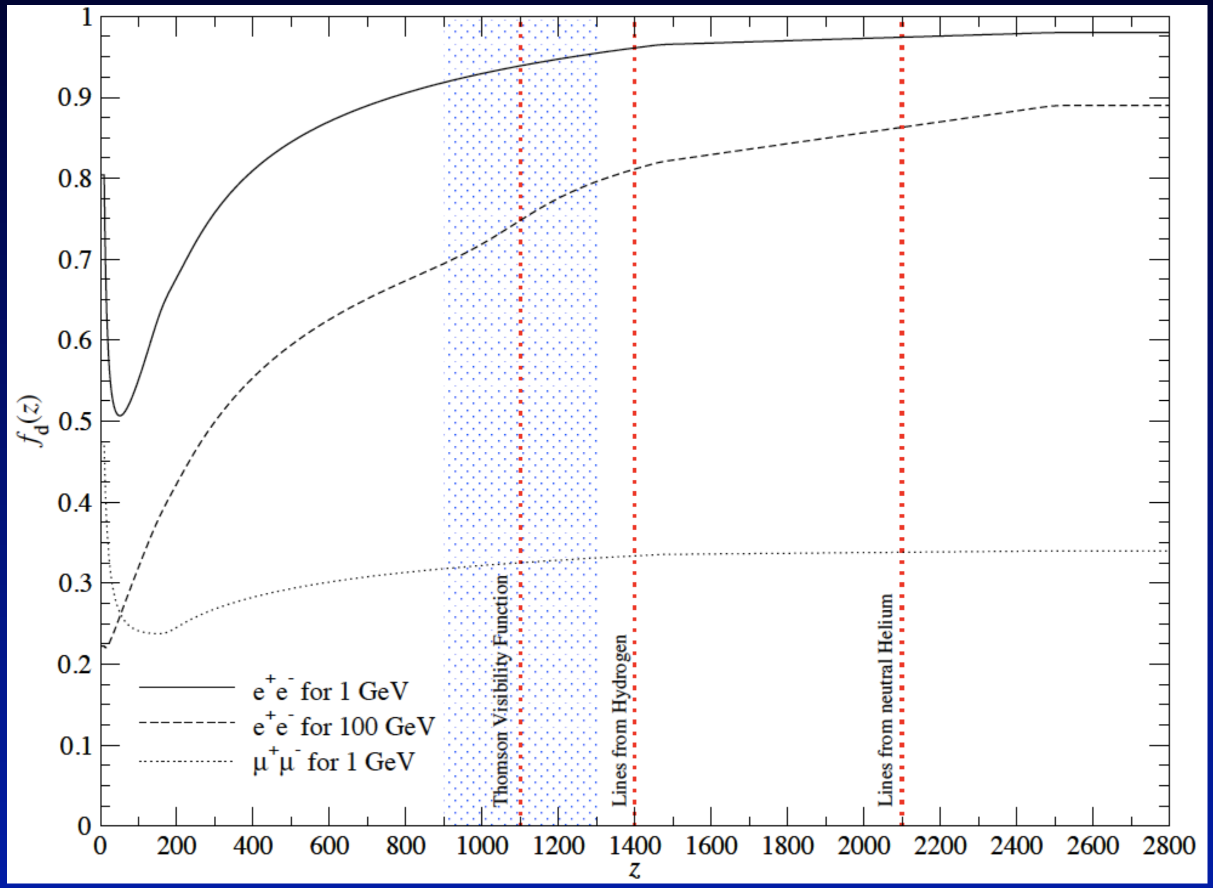


- ‚Hypothetical' source of extra photons parametrized by ε_α & ε_i
- Extra excitations ⇒ delay of Recombination
- Extra ionizations ⇒ affect 'freeze out' tail
- This affects the Thomson visibility function
- From WMAP ⇒ $\varepsilon_\alpha < 0.39$ & $\varepsilon_i < 0.058$ at 95% confidence level (Galli et al. 2008)

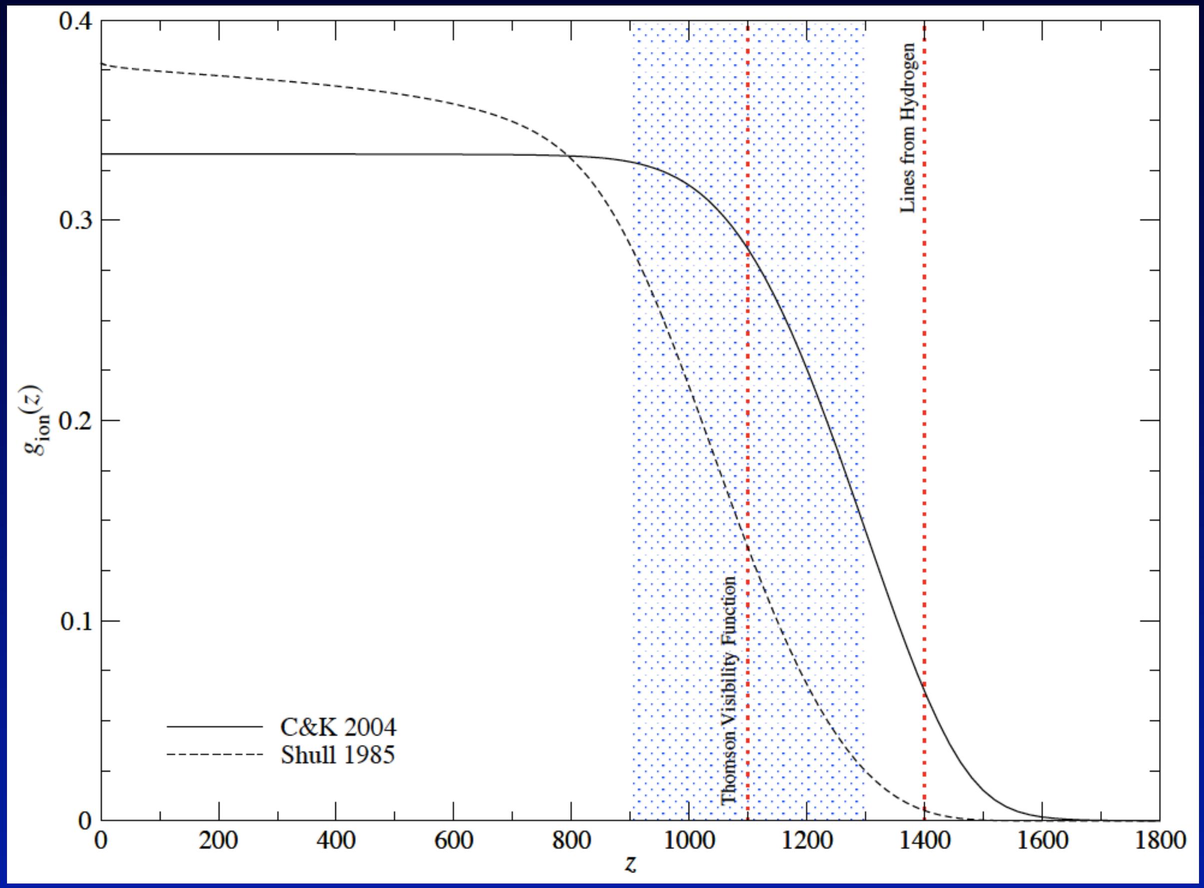
- Extra ionizations & excitations should also lead to additional photons in the recombination radiation!!!
- This in principle should allow us to check for such sources at z~1000

Peebles, Seager & Hu, ApJ, 2000

Dark Matter Annihilation: Energy Branching Ratios



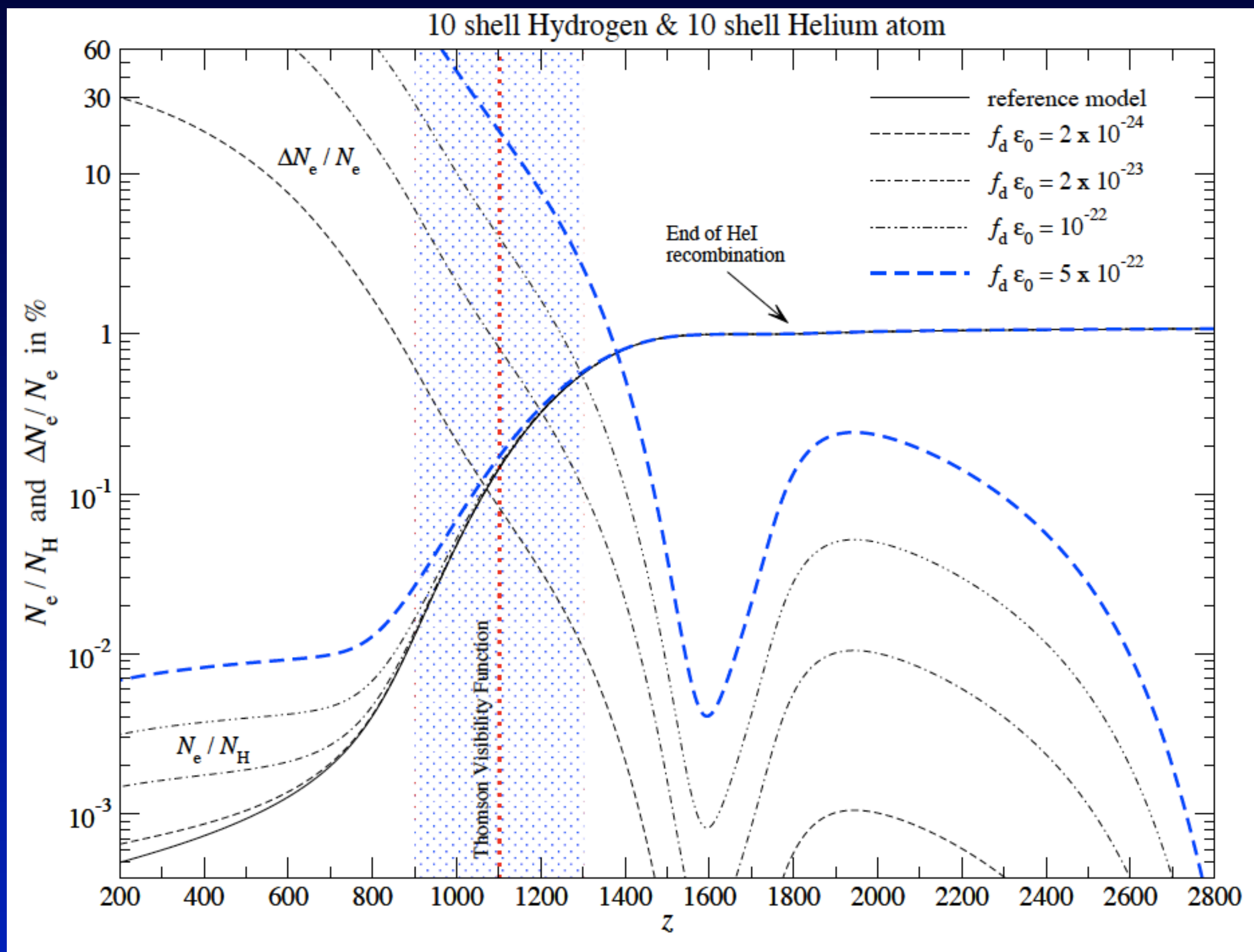
curves from Slatyer et al. 2009

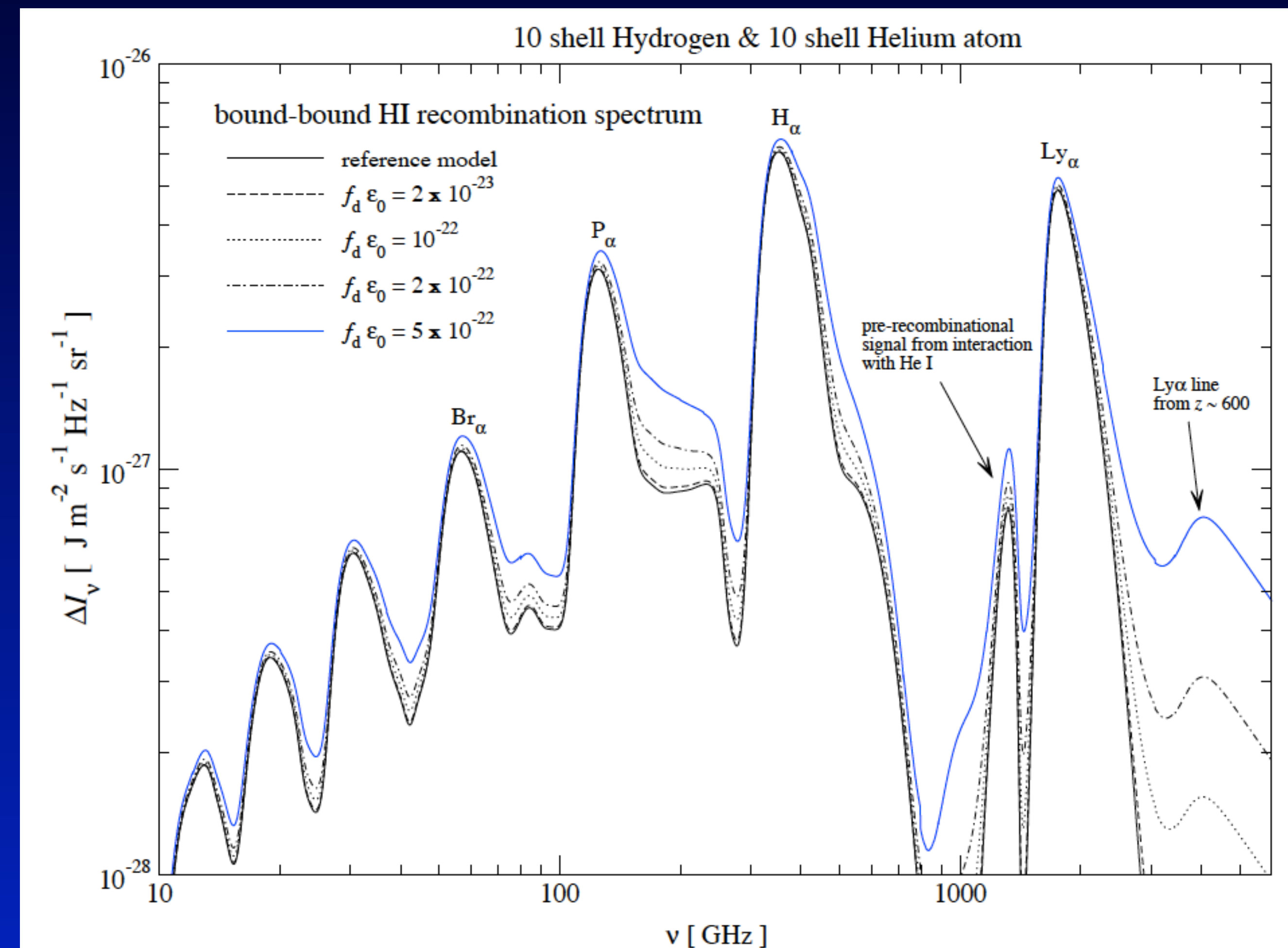


Efficiencies according to Chen & Kamionkowski, 2004 & Shull & van Steenberg 1985

- N^2 - dependence ⇒ dE/dt $\propto (1+z)^6$ and dE/dz $\propto (1+z)^{3...3.5}$
- only part of the energy is really deposited ($f_d \sim 0.1$)
- Branching into *heating* (100% at high z), *ionizations* and *excitations* (mainly during recombination)
- Branching depends on considered DM model

Dark Matter Annihilation: Effect on CMB Anisotropies and the Recombination Spectrum





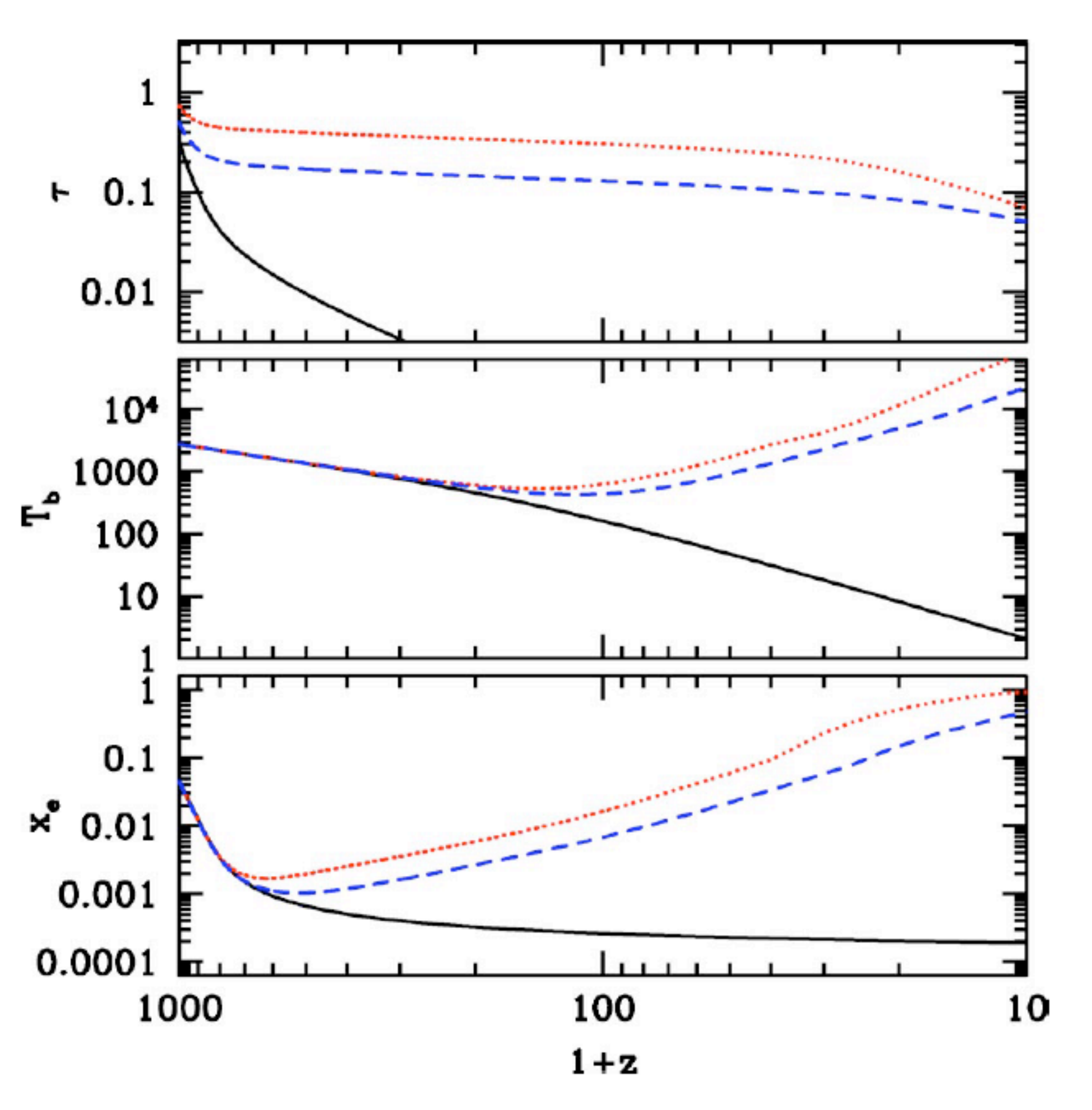
- ‘Delay of recombination’
- Affects Thomson visibility function
- Possibility of Sommerfeld-enhancement
- Clumpiness of matter at z<100

- Additional photons at all frequencies
- Broadening of spectral features
- Shifts in the positions

compare also Padmanabhan & Finkbeiner, Phys. Rev. D, 2005

JC, 2009, arXiv:0910.3663

Decaying particle during & after recombination



Chen & Kamionkowski, 2004

- Modify recombination history
- this changes Thomson visibility function and thus the CMB temperature and polarization power spectra
- ⇒ CMB anisotropies allow

  probing particles with lifetimes ≳ 10^{12} sec
- CMB spectral distortions provide complementary probe! (more tomorrow)