



Stellar Wind Feedback in Massive Young Star Forming Regions

Hazel Rogers and Julian Pittard

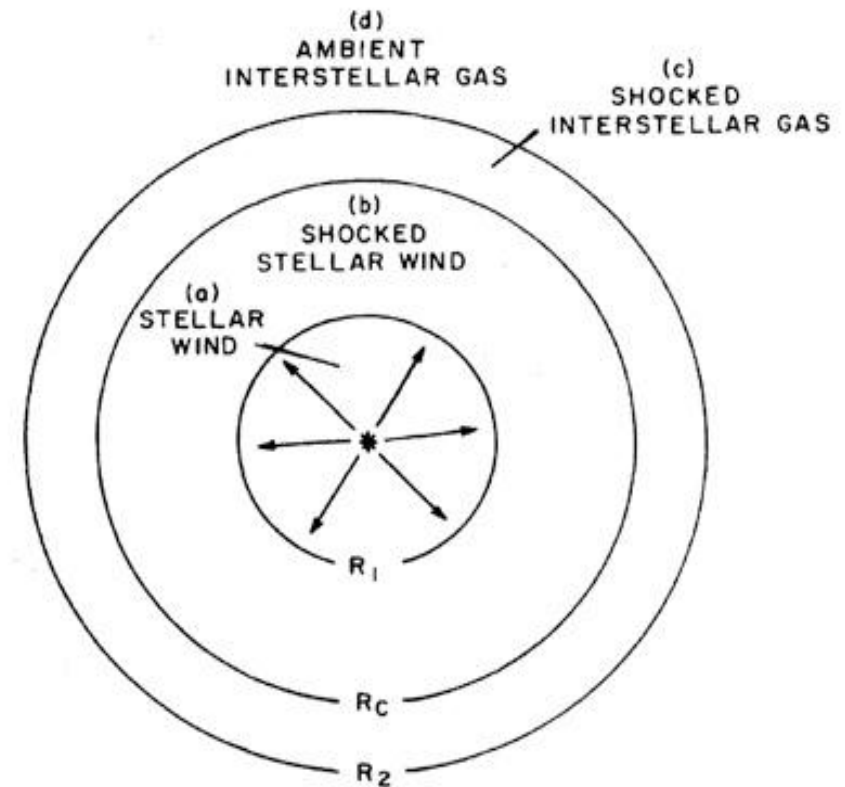


NGC 7635 – The Bubble Nebula

Structure of a bubble:

3 simplified models

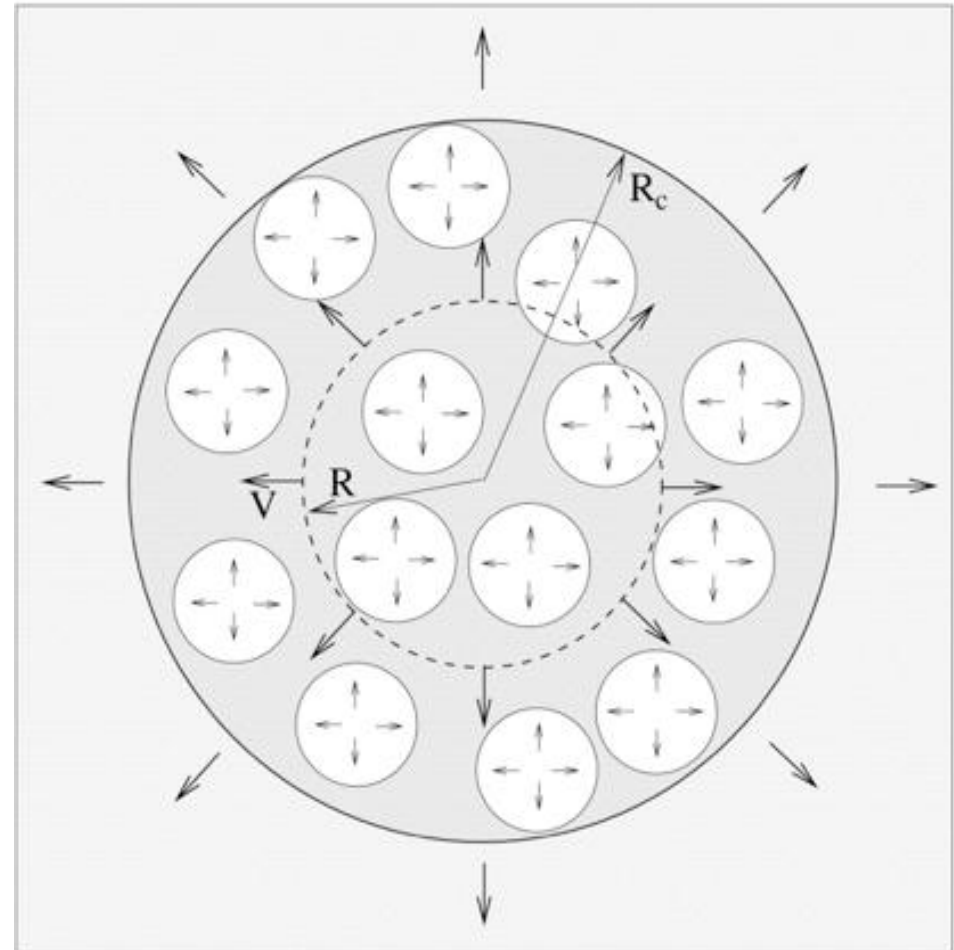
- Castor et al. (1975)
- Chevalier & Clegg (1985)
- Harper-Clark & Murray (2009)



Structure of a bubble:

3 simplified models

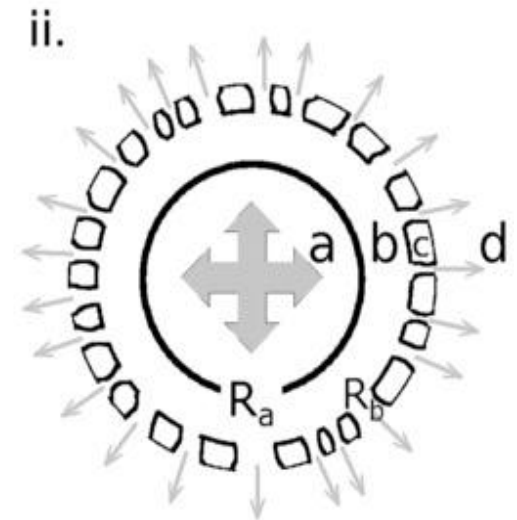
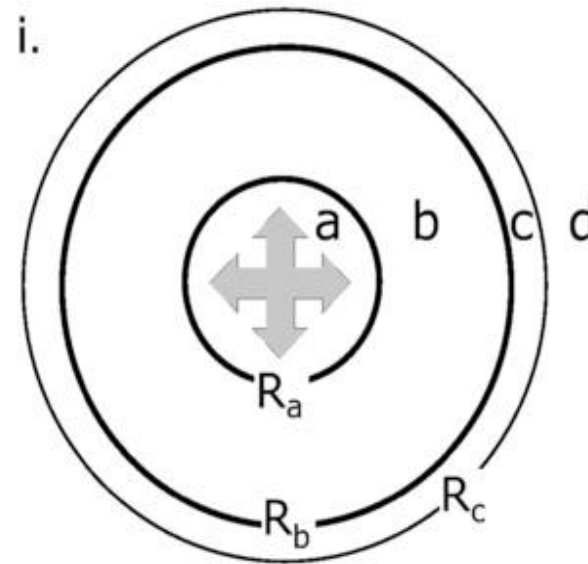
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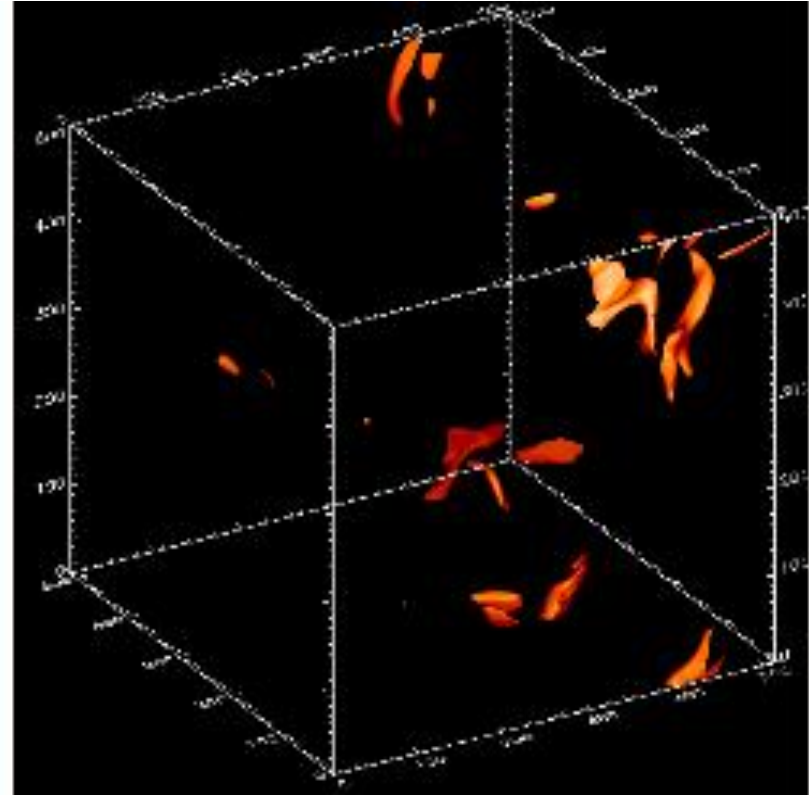
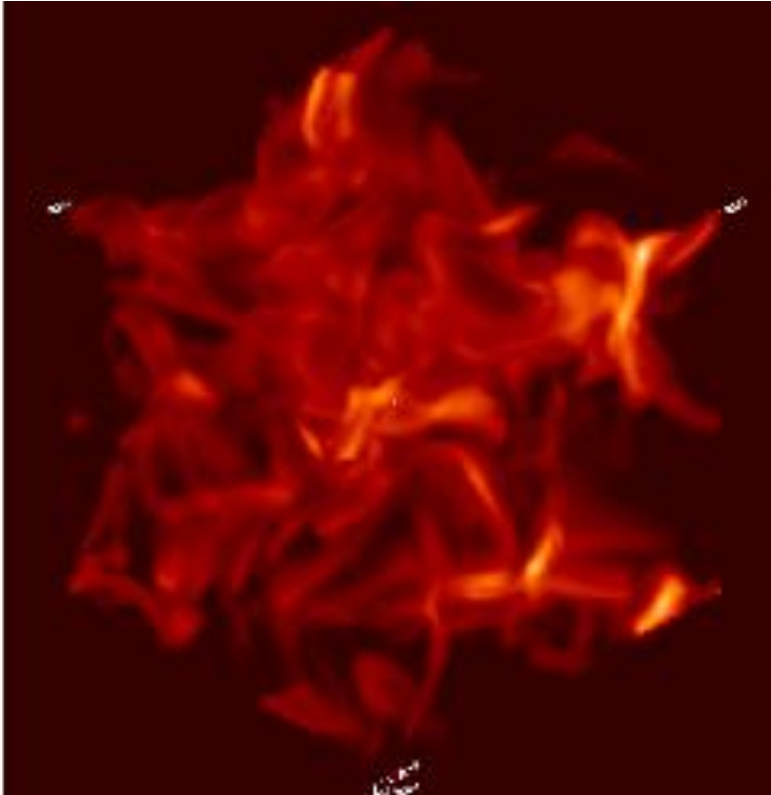
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Simple modelling of stellar winds:

- Used UG code to simulate expanding winds from three massive young O stars
- Stars located at $(0,0,0)$
- Winds expand into inhomogeneous environment
- GMC clump has 5 pc radius
- Total simulation size ± 16 pc
- Heating prescription allows three stable phases



Initial conditions of the GMC based on turbulent ISM simulations of Vasquez-Semadeni et al. (2008)

Evolution of the stars:

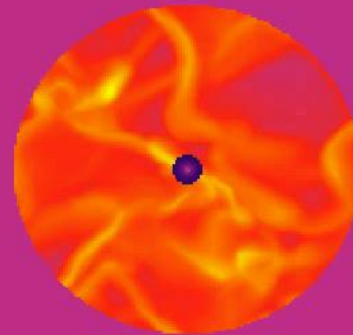
- The three stars evolve over the course of the simulation
 - $30 M_{\odot}$, $25 M_{\odot}$, $20 M_{\odot}$

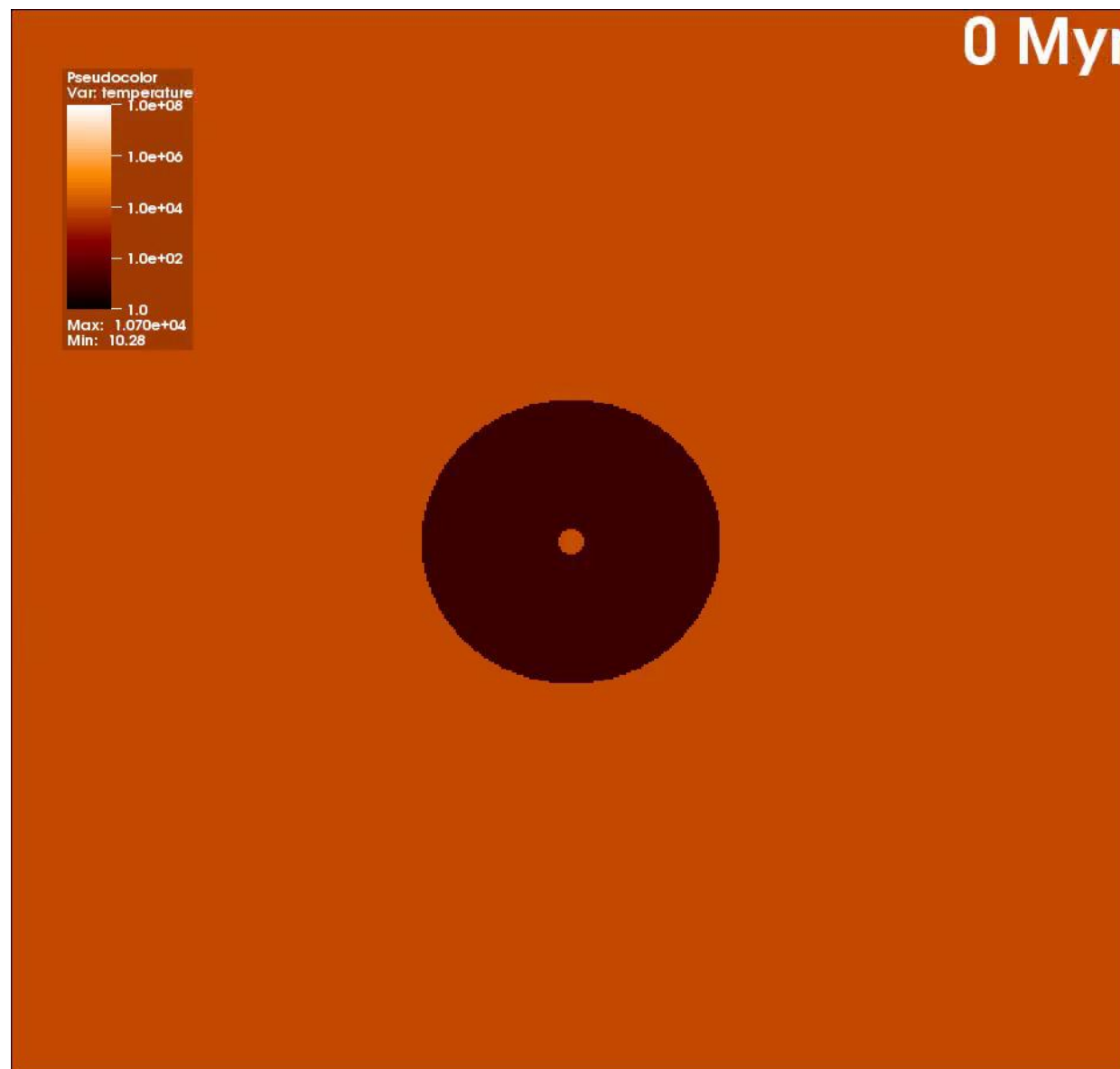
	Mass – loss (M_{\odot}/yr)	V_{inf} (km/s)	Duration (Myrs)
MS	5×10^{-7}	2,000	4
RSG	1×10^{-4}	50	0.1
WR	3×10^{-5}	2,000	0.3

- The $25 M_{\odot}$ star evolves to a RSG shortly after the first Sne
- The $20 M_{\odot}$ star evolves to a RSG shortly after the second Sne



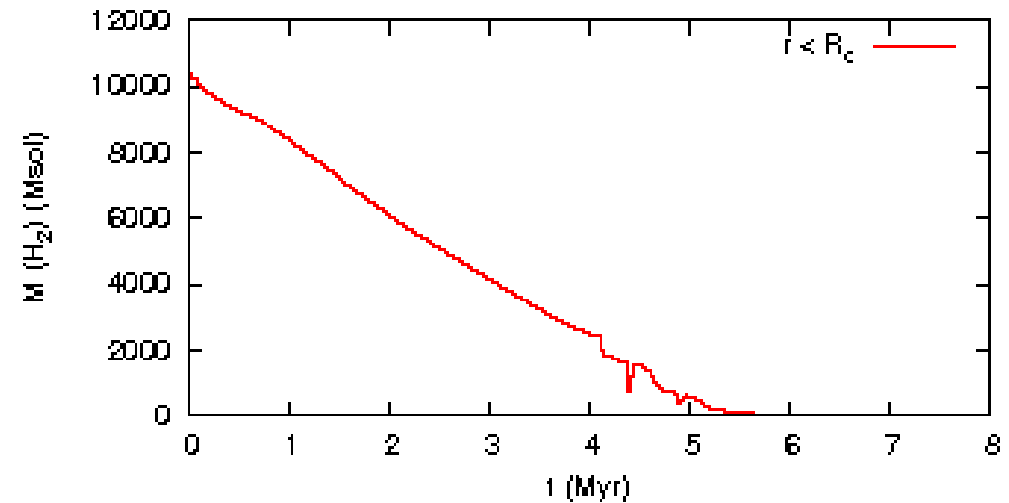
0 Myr





Evolution of H_2 mass of the cluster:

- Winds steadily evacuate clump gas from within the original clump radius
- Evolutionary stages of the stars can be seen
- By the end of the simulation almost all clump gas has been removed from the clump radius





Summary:

- Winds from stars form bubbles in the ISM
- Simulations show the expansion of winds from O stars into an inhomogeneous, turbulent GMC
- The structure of the gas has a large impact on the early expansion of the winds and destruction of the clump
- High density clumps can survive throughout the lifetime of the cluster, although highly ablated
- The supernovae, whilst impressive, do not have as big an effect as perhaps expected